

A New Formalism of the Einstein Equations for Relativistic Rotating Systems

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Combining Geroch's and the ADM formalisms, we give a new formalism for treating the dynamical problems of space-time having a rotational Killing vector. We have found that the basic equations for a rotating system strongly resemble the Einstein and the Maxwell equations for a non-rotating system.

The numerical approach to the relativistic gravitational collapse has been studied by several authors.^{1),2)} However, none of them have succeeded in calculating the collapse of a rotating star though Nakamura, Maeda, Miyama and Sasaki¹⁾ succeeded in constructing the initial data. For a non-rotating star, Nakamura, Maeda, Miyama and Sasaki³⁾ have calculated the formation of black holes. But if one tries to develop their code to a rotating case, one thinks that the basic equations will be terribly complicated. In this short note, we give a new formalism in which the basic equations are simpler and which has some nice features.

In the axially symmetric space-time, there is a rotational Killing vector. This implies we can divide out the Killing direction using Geroch's three-dimensional formalism for space-time with one Killing vector.⁴⁾ Applying this formalism to an axially symmetric system, we get the 3-dimensional "Einstein" equations with respect to $h_{\mu\nu}$ plus some new equations for λ and ω_μ , that is,

$$\begin{aligned} {}^{(3)}R_{\mu\nu} &= (2\lambda^4)^{-1} [\omega_\mu \omega_\nu - h_{\mu\nu} \omega^\rho \omega_\rho] + \lambda^{-1} D_\mu D_\nu \lambda \\ &\quad + 8\pi [\mathcal{T}_{\mu\nu} - (1/2) h_{\mu\nu} (\mathcal{T}^\rho_\rho + \lambda^{-2} \mathcal{T})], \end{aligned} \quad (1)$$

$$\begin{aligned} \lambda^{-1} D^\rho D_\rho \lambda &= - (2\lambda^4)^{-1} \omega^\alpha \omega_\alpha \\ &\quad - 4\pi (\lambda^{-2} \mathcal{T} - \mathcal{T}_\alpha^\alpha), \end{aligned} \quad (2)$$

$$D_{[\mu} \omega_{\nu]} = 8\pi \lambda \varepsilon_{\mu\nu\rho} \mathcal{T}^\rho, \quad (3)$$

$$D^\rho [\lambda^{-3} \omega_\rho] = 0, \quad (4)$$

where

$$\lambda^2 = \xi_\mu \xi^\mu, \quad (5)$$

$$h_{\mu\nu} = g_{\mu\nu} - \lambda^{-2} \xi_\mu \xi_\nu, \quad (6)$$

$$\omega_\mu = \varepsilon_{\mu\nu\rho\sigma} \xi^\nu \nabla^\rho \xi^\sigma, \quad (7)$$

$$\varepsilon_{\mu\nu\rho} = \lambda^{-1} \xi^\sigma \varepsilon_{\mu\nu\rho\sigma},$$

and ξ^μ , ∇^μ , D^μ and ${}^{(3)}R_{\mu\nu}$ represent a rotational Killing vector, the 4-dimensional covariant differentiation, the 3-dimensional covariant differentiation and the 3-dimensional Ricci tensor, respectively.⁴⁾ \mathcal{T} , \mathcal{T}_μ and $\mathcal{T}_{\mu\nu}$ are defined as

$$\mathcal{T} = T_{\mu\nu} \xi^\mu \xi^\nu,$$

$$\mathcal{T}^\rho = h^{\rho\mu} \xi^\nu T_{\mu\nu}$$

and

$$\mathcal{T}_{\rho\sigma} = h_\rho{}^\mu h_\sigma{}^\nu T_{\mu\nu}.$$

We now reduce Eqs. (1)~(4) to the canonical Hamiltonian form applying the ADM formalism.⁵⁾ We first define a new projection tensor*)

*) Small Latin indices refer to the range 0, 1, 2, and capital Latin indices to the range 1, 2.

$$H^{ab} = h^{ab} + n^a n^b, \quad (8)$$

where n_a is a unit normal vector of hypersurface $t = \text{constant}$. Carrying out the projection of all the tensors appeared in Eqs. (1)~(4) by using n_a and H_{ab} , we obtain

$$\begin{aligned} \chi^2 - \chi^{AB} \chi_{AB} + {}^{(2)}R &= 2\lambda^{-1} ({}^{(2)}\Delta \lambda - \chi \kappa) \\ &+ (1/2) \lambda^{-4} (\mathcal{Q}_A \mathcal{Q}^A + \mathcal{Q}_H^2) + 16\pi \rho_H, \end{aligned} \quad (9 \cdot a)$$

$$\begin{aligned} \lambda^{-1} (\lambda \chi_A^B)_{\parallel B} - \lambda^{-1} \partial_A \kappa - \partial_A \chi \\ = 8\pi J_A - (1/2) \lambda^{-4} \mathcal{Q}_H \mathcal{Q}_A, \end{aligned} \quad (9 \cdot b)$$

$$\begin{aligned} \partial_0 \chi_{AB} - \eta^C \chi_{AB\parallel C} + \alpha [{}^{(2)}R_{AB} + \chi \chi_{AB}] \\ - 2\alpha \chi_A^C \chi_{CB} - \alpha \chi_{AB} + (\chi_{AC} \eta_{\parallel B}^C + \chi_{BC} \eta_{\parallel A}^C) \\ - \alpha \lambda^{-1} \chi_{AB} + \alpha \lambda^{-1} \kappa \chi_{AB} - (1/2) \alpha \lambda^{-4} \\ \times [\mathcal{Q}_A \mathcal{Q}_B - H_{AB} (\mathcal{Q}_C \mathcal{Q}^C - \mathcal{Q}_H^2)] \\ - 8\pi \alpha [S_{AB} + (1/2) H_{AB} \\ \times (\rho_H - S_C^C - \lambda^{-2} \mathcal{T})], \end{aligned} \quad (10 \cdot a)$$

$$\partial_0 H_{AB} = -2\alpha \chi_{AB} + \eta_{A\parallel B} + \eta_{B\parallel A}, \quad (10 \cdot b)$$

$$\partial_0 \lambda - \eta^A \partial_A \lambda = -\alpha \kappa, \quad (11 \cdot a)$$

$$\begin{aligned} \partial_0 \kappa - \eta^A \partial_A \kappa \\ = \alpha \chi \kappa - H^{AB} (\partial_A \alpha) (\partial_B \lambda) - \alpha {}^{(2)}\Delta \lambda \\ - (1/2) \alpha \lambda^{-3} (\mathcal{Q}_A \mathcal{Q}^A - \mathcal{Q}_H^2) - 4\pi \alpha \lambda \\ \times [\rho_H - S_A^A + \lambda^{-2} \mathcal{T}], \end{aligned} \quad (11 \cdot b)$$

$$\mathcal{Q}_{A\parallel B} - \mathcal{Q}_{B\parallel A} = -16\pi \lambda \varepsilon_{AB} J_\phi, \quad (12 \cdot a)$$

$$\begin{aligned} \partial_0 \mathcal{Q}_A - \eta^B \mathcal{Q}_{A\parallel B} = \eta_{\parallel A}^B \mathcal{Q}_B + \partial_A (\alpha \mathcal{Q}_H) \\ + 16\pi \varepsilon_{AB} \alpha \lambda S^B, \end{aligned} \quad (12 \cdot b)$$

$$\begin{aligned} \partial_0 \mathcal{Q}_H - \eta^A \partial_A \mathcal{Q}_H = (\alpha \mathcal{Q}^A)_{\parallel A} + \alpha \chi \mathcal{Q}_H \\ - 3\alpha \lambda^{-1} [\mathcal{Q}_H \kappa + \mathcal{Q}^A \partial_A \lambda], \end{aligned} \quad (12 \cdot c)$$

where α and η_A are defined by

$$\begin{aligned} ds^2 = h_{ab} dx^a dx^b = -\alpha^2 dt^2 + H_{AB} \\ \times (dx^A + \eta^A dt) (dx^B + \eta^B dt), \end{aligned}$$

and χ_{AB} are the extrinsic curvatures with respect to H_{AB} , and

$$\chi = \chi^A A, \quad \kappa = -n^\mu \partial_\mu \lambda,$$

$$\mathcal{Q}_A = H_A^b \omega_b, \quad \mathcal{Q}_H = n_a \omega^a,$$

$$\varepsilon_{ab} = n^c \varepsilon_{cab},$$

$$\rho_H = n_a n_b \mathcal{T}^{ab}, \quad J_\phi = -n_a \mathcal{T}^a,$$

$$J^A = -n_a H_b^A \mathcal{T}^{ab}, \quad S^A = H_b^A \mathcal{T}^b$$

and

$$S_{AB} = H_{Aa} H_{Bb} \mathcal{T}^{ab}.$$

In Eqs. (9)~(12), ${}^{(2)}\Delta$, \parallel , ${}^{(2)}R$ and ${}^{(2)}R_{AB}$ are the Laplacian, the covariant differentiation, the scalar curvature and the Ricci tensor with respect to H_{AB} , respectively. Equations (9) and (10) are the constraint equations and the evolution equations of two-dimensional space metric, reduced from the 3-dimensional ‘‘Einstein’’ equations. Equations (11) and (12) resemble the Maxwell equations if we consider $(1/2)\xi_\mu$ as the vector potential, that is, if we define the ‘‘electromagnetic’’ field as

$$B_\phi = \mathcal{Q}_H/2, \quad B^A = \varepsilon^{AB} \partial_B \lambda,$$

$$E_\phi = \kappa \lambda \quad \text{and} \quad E^A = \varepsilon^{AC} \mathcal{Q}_C/2\lambda.$$

Using these ‘‘electromagnetic’’ fields, we have the hydrodynamics equations as follows:*)

(i) the energy equation

$$\begin{aligned} \partial_0 (\lambda \sqrt{H} \rho_H) + \partial_B (U^B \lambda \sqrt{H} \rho_H) \\ = -\partial_B (\alpha \lambda \sqrt{H} p V^B) \\ + \alpha \lambda \sqrt{H} p (\chi + \lambda^{-1} \kappa) \\ + \alpha \lambda \sqrt{H} (p + \rho_H) \{-a_B V^B + V^A V^B \chi_{AB}\} \\ + \alpha \lambda \sqrt{H} \lambda^{-2} J_\phi (\lambda^{-2} E_\phi V_\phi + 2E_B V^B), \end{aligned} \quad (13)$$

(ii) the conservation of angular momentum

$$\partial_0 (\lambda \sqrt{H} J_\phi) + \partial_B (U^B \lambda \sqrt{H} J_\phi) = 0, \quad (14)$$

(iii) the Euler equation

$$\begin{aligned} \partial_0 (\lambda \sqrt{H} J_A) + \partial_B (U^B \lambda \sqrt{H} J_A) \\ = -\alpha \lambda \sqrt{H} [\partial_A p + (p + \rho_H) a_A] \\ + \alpha \lambda \sqrt{H} (p + \rho_H) [(1/2) (\partial_A H_{BC}) \\ \times V^B V^C + \alpha^{-1} V_C \partial_A \eta^C] \end{aligned}$$

*) The method of deriving the hydrodynamics equations is the same as Ref. 1).

$$+ \alpha \lambda \sqrt{H} \lambda^{-2} J_\phi [2E_A + \varepsilon_{AC} \\ \times (2\lambda^{-1} B_\phi V^C - \lambda^{-1} B^C V_\phi)] \quad (15)$$

and

(iv) the equation of continuity

$$\partial_0 \left(\frac{\lambda \sqrt{H} n}{\sqrt{1-V^2}} \right) + \partial_B \left(U^B \frac{\lambda \sqrt{H} n}{\sqrt{1-V^2}} \right) = 0, \quad (16)$$

where

$$H = \det(H_{AB}), \quad a_B = \partial_B \ln \alpha,$$

$$V^B = (p + \rho_H)^{-1} J^B, \quad V_\phi = (p + \rho_H)^{-1} J_\phi,$$

$$V^2 = V_B V^B + \lambda^{-2} V_\phi^2 \quad \text{and} \quad U^A = \alpha V^A - \eta^A.$$

p and n are the pressure and the proper number density of matter, respectively.

We can see that the "electromagnetic" fields give Joule's heat-like and the Lorentz force-like terms into the hydrodynamics equations. Note that the angular momentum density behaves like charge both in Eqs. (12) and (14).

We want to call the present new formalism the $[(2+1)+1]$ -dimensional representation of the Einstein equations. On the other hand the ADM formalism is called the $(3+1)$ -dimensional representation of the Einstein equations. Comparing these two for a rotating system, we find that the $[(2+1)+1]$ -dimensional formalism has the following merits:

(1) Since the ϕ -component of the shift

vector does not appear in the basic equations, we do not need to consider the coordinate condition about ϕ -component.

(2) In the $(3+1)$ -dimensional formalism, for a rotating case we should treat $g_{A\phi}$ and $\dot{g}_{A\phi}$ whose behaviours are unknown. In the $[(2+1)+1]$ -dimensional formalism, new variables which are needed for a rotating case can be considered the electromagnetic fields. So we can understand their behaviours easily. These facts will enable us to utilize our code for a non-rotating collapse almost as it is.

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